

MODELING THE INFRARED EMISSION IN CYGNUS A

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Received 2011 October 18; accepted 2011 December 13; published 2012 February 13

ABSTRACT

We present new *Spitzer* IRS spectroscopy of Cygnus A, one of the most luminous radio sources in the local universe. Data on the inner 20'' are combined with new reductions of MIPS and IRAC photometry as well as data from the literature to form a radio through mid-infrared spectral energy distribution (SED). This SED is then modeled as a combination of torus reprocessed active galactic nucleus (AGN) radiation, dust enshrouded starburst, and a synchrotron jet. This combination of physically motivated components successfully reproduces the observed emission over almost 5 dex in frequency. The bolometric AGN luminosity is found to be $10^{12} L_{\odot}$ (90% of L_{IR}), with a clumpy AGN-heated dust medium extending to ~ 130 pc from the supermassive black hole. Evidence is seen for a break or cutoff in the core synchrotron emission. The associated population of relativistic electrons could in principle be responsible for some of the observed X-ray emission though the synchrotron self-Compton mechanism. The SED requires a cool dust component, consistent with dust-reprocessed radiation from ongoing star formation. Star formation contributes at least $6 \times 10^{10} L_{\odot}$ to the bolometric output of Cygnus A, corresponding to a star formation rate of $\sim 10 M_{\odot} \text{ yr}^{-1}$.

Key words: galaxies: active – galaxies: individual (Cygnus A) – galaxies: jets

Online-only material: color figures

1. INTRODUCTION

As one of the nearest powerful radio-loud active galactic nuclei (AGNs), Cygnus A provides an excellent laboratory to study the environment and activity of powerful AGNs. The luminosity of the AGN in Cygnus A (which is predominantly expressed in the infrared) is high enough to classify it as a quasar (e.g., Djorgovski et al. 1991). While not seen in total intensity, broad H α is seen in polarized light (Ogle et al. 1997), lending support for the existence of an obscured broad-line region (BLR). The polarized broad lines were detected within the ionization cone seen by Jackson et al. (1996), giving them a possible scattering origin in the narrow line region (NLR).

The host galaxy of Cygnus A is a cD elliptical near the center of a cluster which appears to be undergoing a merger with a cluster of similar size (Ledlow et al. 2005). Tadhunter et al. (2003) obtained *Hubble Space Telescope* (HST) Space Telescope Imaging Spectrograph (STIS) spectra of the nuclear region. Stepping the slit across the nucleus, a velocity gradient indicative of rotation around the radio axis was observed. Modeling the velocity as due to the potential of a supermassive black hole (SMBH) and stellar mass distribution (measured from a $1.6 \mu\text{m}$ NICMOS image) gives an SMBH mass of $(2.5 \pm 0.7) \times 10^9 M_{\odot}$. This measurement is consistent with black-hole-mass–host-galaxy relations (e.g., Magorrian et al. 1998; Ferrarese & Merritt 2000; Gebhardt et al. 2000; Gültekin et al. 2009).

The large-scale radio morphology shows prominent hot spots, lobes, and a radio core. Both a jet and counterjet are visible in Very Large Array (VLA) and Very Long Baseline Array (VLBA) observations (e.g., Sorathia et al. 1996). Due to the

edge-brightened morphology it is classified as an FR II source (Fanaroff–Riley Type II; Fanaroff & Riley 1974).

X-ray observations of Cygnus A are consistent with the presence of a hidden quasar. *Chandra* ACIS observations from 0.7 to 9 keV by Young et al. (2002) show the hard X-ray flux to be peaked at a location consistent with that of the radio core and unresolved (less than 0''.4 in size, determined by comparison with a model point-spread function). Higher energy INTEGRAL observations show emission between 20 and 100 keV (Beckmann et al. 2006), although there may be some contamination from intracluster gas. This hard X-ray emission is likely due to accretion disk emission Comptonized in the AGN corona, although the UV/optical emission is obscured.

In addition to the AGN activity, *HST* imaging has also revealed star formation in the central region of Cygnus A, which began < 1 Gyr ago (Jackson et al. 1998). It is located in a 4 kpc ring around the nucleus, oriented orthogonal to the radio axis.

Based on adaptive optics observations showing a secondary point source near the nucleus, Canalizo et al. (2003) suggest that Cygnus A may be in the late stages of a merger event. This merger event and related accretion may be related to the current epoch of nuclear activity.

A near-infrared spectrum presented by Bellamy & Tadhunter (2004) showed complicated emission line properties suggesting an infalling molecular cloud, consistent with the Canalizo et al. (2003) picture of a minor merger. The H $_2$ lines were seen in several components, both redshifted and blueshifted relative to the systemic velocity, interpreted as emission from a rotating torus. The observed near-infrared line ratios are consistent with excitation by X-rays (from the AGN), while likely ruling out shocks as a possible excitation method.

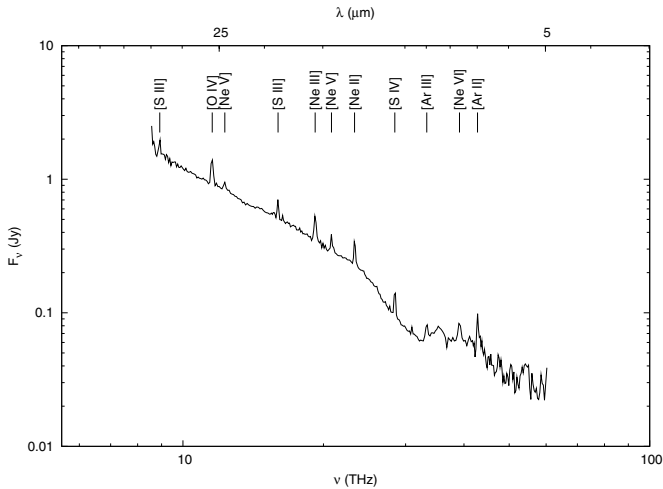


Figure 1. *Spitzer* IRS spectrum for Cygnus A. Extracted in a 20'' circular aperture from spectral mapping data.

For a more detailed summary of Cygnus A, including properties of the larger host galaxy and environment, see the review by Carilli & Barthel (1996).

While the bulk of the infrared emission is likely related to the presence of an AGN, the star formation can contribute to the infrared emission as well. Additionally, the bolometric luminosity of the AGN is uncertain. Estimates have been made using template spectral energy distributions (SEDs) and X-ray observations. However, a significant portion of the bolometric luminosity comes out in the infrared, via dust reprocessing of the UV/optical continuum. An accurate determination of the bolometric luminosity then requires an understanding of the contributions to the observed infrared luminosity.

In this paper, we present modeling of the SED of Cygnus A. To accomplish this we combined new infrared measurements obtained by the *Spitzer Space Telescope*'s Infrared Spectrograph (IRS; Werner et al. 2004) with existing *Spitzer* observations with the Imaging Array Camera (IRAC; Fazio et al. 2004) and the Multiband Imaging Photometer for SIRTf (MIPS; Rieke et al. 2004), and measurements of the radio core from the literature to construct a radio through infrared ($\sim 2\text{--}10^5$ GHz; $4\text{--}10^5$ μm) SED of the inner regions of Cygnus A. The resulting SED was subsequently modeled using components intended to replicate the physical processes likely to produce the observed emission. Using the results of the fitting we have been able to decompose the infrared emission and determine the bolometric luminosity of the AGN in Cygnus A. This in turn also provides an estimate of the star formation rate.

The paper is organized as follows: in Section 2, we discuss the new *Spitzer* IRS observations, new reductions of *Spitzer* IRAC and MIPS data, as well as the data from the literature. In Section 3, we describe the components used to model the SED, for which the results are described in Section 4. We conclude with a general discussion in Section 5 and a summary in Section 6.

2. DATA

New observations were combined with data compiled from multiple archives and new reductions in order to obtain an SED covering almost five orders of magnitude in frequency. Our analysis is focused on the properties of the continuum emission in this system. Below we discuss the reduction and analysis

Table 1
Mid-Infrared Emission Line Properties

Line	Flux ($\times 10^{-13}$ erg s $^{-1}$ cm $^{-2}$)	FWHM (μm)
[Ar II] 6.9	1.93 ± 0.08	0.11
[Ne VI] 7.6	1.37 ± 0.09	0.11
[Ar III] 8.9	1.11 ± 0.04	0.11
[S IV] 10.5	1.55 ± 0.03	0.09
[Ne II] 12.8	2.67 ± 0.03	0.11
[Ne V] 14.3	1.51 ± 0.04	0.10
[Ne III] 15.6	4.13 ± 0.04	0.15
[S III] 18.7	2.42 ± 0.06	0.13
[Ne V] 24	2.51 ± 0.06	0.37
[O IV] 25.9	4.58 ± 0.08	0.32
[S III] 33	3.33 ± 0.08	0.31
H ₂ S(7)	0.944 ± 0.130	0.06
H ₂ S(5)	0.853 ± 0.112	0.06
H ₂ S(3)	0.286 ± 0.030	0.09
H ₂ S(2)	0.333 ± 0.025	0.11
H ₂ S(0)	0.128 ± 0.038	0.31

Notes. Only detections $>3\sigma$ are listed. Errors quoted from PAHFIT output.

of the new *Spitzer* IRS observations and the re-reduction of archival IRAC and MIPS observations

2.1. *Spitzer* IRS Observations

Cygnus A was observed (PI: Baum) in “mapping mode” using the low-resolution mode of IRS on board the *Spitzer Space Telescope*. Both the short- and long-wavelength slits were stepped across the source in increments of half the slit width. After pipeline calibration at the *Spitzer* Science Center, the observations were combined into a spectral data cube using the Cube Builder for IRS Spectra Maps (CUBISM; Smith et al. 2007a). From this data cube, a spectrum was extracted in a 20'' aperture (Figure 1).

2.1.1. Removal of Mid-Infrared Emission Lines

The IRS spectrum was fit using PAHFIT (Smith et al. 2007b) to measure and remove contributions from narrow emission lines. Integrated line fluxes and widths are provided in Table 1. While a study of the emission line properties is beyond the scope of this paper, we note the detection of multiple high ionization lines such as [O IV], [Ne V], and [Ne VI]. These are all consistent with the presence of an AGN (Genzel et al. 1998; Armus et al. 2007, and references therein). Of particular note are [Ne V] and [Ne VI] which require the presence of ionizing photons of at least 97.1 and 125.8 eV, respectively. All the measured emission lines in Table 1 were subtracted from the IRS spectrum before fitting the SED.

2.1.2. Dust Features

In addition to fitting the emission lines, we have also fit the mid-infrared dust features using PAHFIT. Table 2 shows the integrated flux and profile FWHM based on the PAHFIT output. In contrast with the emission lines, the dust features were not removed as they are expressed in the starburst models.

Spoon et al. (2007) developed a diagnostic diagram using the $6.2\text{ }\mu\text{m}$ polycyclic aromatic hydrocarbon (PAH) and the strength of the $9.7\text{ }\mu\text{m}$ silicate feature (S_{sil}), which is seen in absorption in Cygnus A. These two spectral features can be used in tandem to classify the relative dominance of PAH emission, continuum

Table 2
Mid-Infrared Dust Features

λ (μm)	Flux ($\times 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2}$)	FWHM (μm)
6.2	1.76 ± 0.23	0.19
6.7	5.99 ± 0.49	0.47
7.4	22.3 ± 0.8	0.94
7.8	2.26 ± 0.32	0.42
8.3	5.50 ± 0.23	0.42
8.6	4.42 ± 0.20	0.34
11.3	0.685 ± 0.101	0.36
12.0	2.71 ± 0.12	0.54
12.6	3.82 ± 0.21	0.53
13.5	2.49 ± 0.12	0.54
14.2	1.28 ± 0.11	0.36
17.0	2.52 ± 0.27	1.11
17.4	0.397 ± 0.065	0.21
17.9	0.802 ± 0.099	0.29
18.9	2.19 ± 0.14	0.36
33.1	3.54 ± 0.51	1.66

Notes. Only detections $>3\sigma$ are listed. Errors quoted from PAHFIT output.

Table 3
Additional Infrared Flux Densities from *Spitzer*

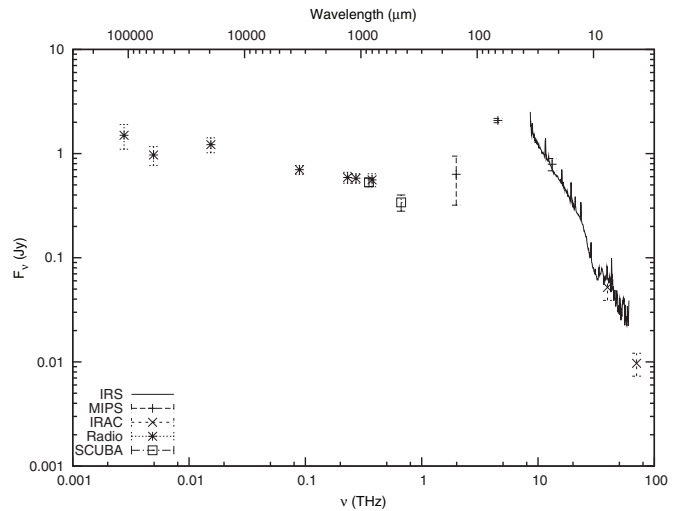
Instrument/Channel	λ (μm)	F_ν (Jy)	σ (Jy)	Aperture ($''$)
IRAC-2	4.5	0.010	0.003	12.2
IRAC-4	8	0.054	0.013	12.2
MIPS-70	70	2.20	0.11	30
MIPS-160	160	0.668	0.033	48

emission, and silicate absorption. We follow their method of spline fitting to measure the depth of the silicate feature, finding $S_{\text{sil}} \approx -0.8$, where S_{sil} is the negative of the apparent optical depth. The equivalent width of the $6.2 \mu\text{m}$ PAH is $0.0552 \mu\text{m}$, placing Cygnus A on the border of regions 1A and 2A in their diagnostic diagram, corresponding to objects dominated by continuum emission in the mid-infrared. This is consistent with the presence of a strong AGN.

2.2. Archival *Spitzer* IRAC + MIPS Data

The nucleus and hot spots were imaged using IRAC at 4.5 and $8.0 \mu\text{m}$ (see Stawarz et al. 2007, for a presentation of the data and analysis of hot spot properties). As fluxes for the core were not presented, the archival data were retrieved, re-reduced, and calibrated according to the IRAC instrument manual. The images showed strong emission at the location of both radio hot spots and the radio core in both channels. The emission was unresolved in the core component, the measured flux densities for this component are given in Table 3. Extraction apertures of $12''/2$ were used for both channels.

Shi et al. (2005) presented observation of Cygnus A using the MIPS instrument on *Spitzer* at 24, 70, and $160 \mu\text{m}$. Their $24 \mu\text{m}$ flux is consistent with our IRS observations. The slope between their 70 and $160 \mu\text{m}$ points is steeper than that of the Rayleigh–Jeans tail. The 70 and $160 \mu\text{m}$ data were re-reduced from the BCD products in the *Spitzer* Science Center archive. Our measured flux densities are given in Table 3. The core of Cygnus A was unresolved in both the 70 and $160 \mu\text{m}$ bands. Apertures of $30''$ and $48''$ were used at 70 and $160 \mu\text{m}$, respectively. Our $160 \mu\text{m}$ measurement has larger error bars, but is broadly consistent with their value.

**Figure 2.** Radio through mid-infrared SED with new *Spitzer* observations and other data from the literature. See the text for details and references.**Table 4**
Flux Densities from the Literature

λ (μm)	F_ν (Jy)	σ (Jy)	Resolution ($''$)	References
450	0.34	0.06	8	1
800	0.56	0.08	13	2
850	0.53	0.05	14	1
1100	0.58	0.06	19	2
1300	0.59	0.07	11	3
3.3×10^3 (89 GHz)	0.70	0.07	2	4
19.5×10^3 (15.4 GHz)	1.22	0.20	Not provided in ref.	5
60.6×10^3 (5 GHz)	0.97	0.20	2.0×3.1	5
109×10^3 (2.7 GHz)	1.5	0.4	3.7×5.8	5

References. (1) Robson et al. 1998; (2) Eales et al. 1989; (3) Salter et al. 1989; (4) Wright & Birkinshaw 1984; (5) Alexander et al. 1984.

2.3. Published Data

Cygnus A's strong synchrotron emission at radio frequencies suggests that this may contribute to the infrared as well. To anchor the synchrotron spectrum we supplemented the *Spitzer* observations with radio measurements from the literature. Core fluxes were obtained from Eales et al. (1989), Salter et al. (1989), Alexander et al. (1984), and Wright & Birkinshaw (1984). Submillimeter nuclear fluxes were also taken from Robson et al. (1998). Based on these archival data the unresolved radio core⁹ is flat spectrum ($\alpha = 0.18$, $F_\nu \propto \nu^{-\alpha}$), until ~ 1 THz, where thermal emission from dust begins to dominate the SED. The compiled radio through mid-infrared SED is shown in Figure 2 and the flux densities taken from the literature are provided in Table 4. The resolution of the observations is also provided.

There is some uncertainty associated with the highest frequency sub-mm observations. The $450 \mu\text{m}$ observation suggests that the synchrotron break may be occurring in this spectral regime. It is unclear if this is a genuine break or if the measurements are affected by variability. Additional observations at these frequencies may be able to clarify this issue.

A consideration of the apertures is important when assembling data across several orders of magnitude in frequency. The

⁹ Here “radio core” refers to the core seen by the VLA with on kpc-scale resolution. This encompasses flux from the very long baseline interferometry (VLBI) scale core and jet.

resolution of the data used to assemble the SED in Figure 2 varies by over a factor of 10, but in all cases the core component is unresolved. At lower frequencies the emission is due solely to synchrotron emission from the flat spectrum radio core. There is no evidence to suggest that the larger-scale steep spectrum jet will contribute emission in the infrared. Accordingly we use the unresolved VLA-scale core to anchor the synchrotron emission in the nucleus, so the difference in apertures should not affect the construction of this SED for the nucleus.

3. MODELING

We aim to reproduce the major features in the SED: the power-law emission at radio wavelengths, the strong thermal emission at infrared wavelengths, and the overall character of the silicate absorption. Extrapolating the power law from the radio to the infrared requires a modification of the power-law spectrum to avoid exceeding the observed mid-infrared flux. The infrared emission is thermal in nature, coming from dust at a variety of temperatures ranging from the cold interstellar medium (ISM) ($T \sim 20$ K) through hot dust near the AGN up to the sublimation temperature ($T \sim 1500$ K). The power source for this dust heating is a combination of star formation and AGN activity, with an uncertain balance between the two.

As noted in the introduction, previous studies of Cygnus A have found evidence for simultaneous ongoing AGN activity and star formation, both of which can contribute to the infrared emission. We model the continuum emission from ~ 2 to 10^5 GHz (3×10^6 – $5 \mu\text{m}$; see Figure 2) using components to represent the AGN torus model, a starburst, and synchrotron emission. The selection of models has 15 free parameters, and the particular choices are discussed in the following subsections.

3.1. AGN Torus Model

According to the unified scheme for radio-loud AGNs, the SMBH and accretion disk can be hidden from view along some lines of sight by an obscuring torus. Along these obscured lines of sight the UV/optical radiation is absorbed and re-radiated in the mid- and far-infrared. As noted above, Cygnus A shows evidence for a hidden BLR through observations of $H\alpha$ in polarized light. This suggests that Cygnus A harbors a “hidden” accretion disk and a BLR which is obscured along our line of sight.

Nenkova et al. (2008) have constructed a model for an obscuring AGN torus, where the obscuration is due to the presence of multiple clouds along the line of sight to the AGN. We select this set of models because clumpy models seem to provide better fits to Sil features than smooth dust distributions (e.g., Baum et al. 2010). In order to reproduce the observed ratio of Type I/II AGNs, the obscuring clouds collectively populate a rough toroidal structure with some opening angle.

The CLUMPY torus model is specified by multiple geometrical parameters. The outer radius of the torus (R_o) is Y times the inner radius (R_d), where the inner radius is determined from a dust sublimation temperature of $T = 1500$ K ($R_d = 0.4 \times L_{45}^{0.5}$ pc, L_{45} is the bolometric luminosity of the AGN in units of 10^{45} erg s $^{-1}$). The geometry of the model is shown in Figure 3. The clumps have a Gaussian angular distribution, with σ parameterizing the width of the angular distribution from the mid-plane. The radial distribution is a power law with index q : r^{-q} . The inclination of the torus symmetry axis to the line of sight is i and the average number of clouds along a given line of sight is N .

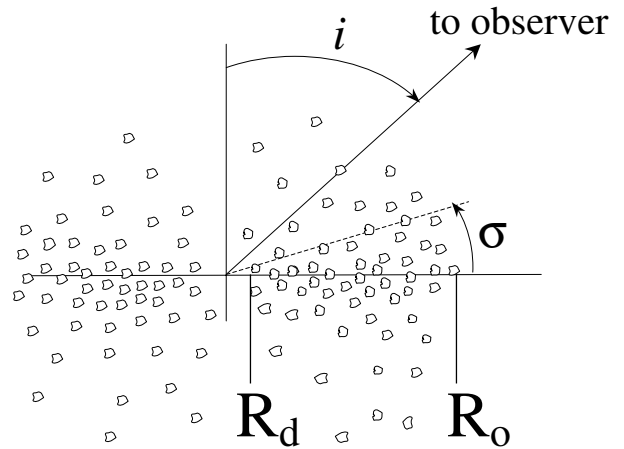


Figure 3. Geometry of the CLUMPY torus model. Figure from Nenkova et al. (2008), used with permission. See the text for an explanation of the labels.

In addition to the geometrical parameters, the models also vary the optical depth of each clump (τ_V), using the Ossenkopf et al. (1992) dust composition and Mathis et al. (1977) grain size distribution. The overall scaling of the model flux is F_{AGN} , the bolometric flux of the AGN’s accretion disk component (treating the synchrotron emission separately).

The clumpy nature of the obscuration implies that there is a finite probability of a direct line of sight to the BLR, even at high inclinations. However, as Cygnus A shows no evidence for directly observed broad lines, we have only fit models where the central regions are fully obscured along our line of sight.

3.2. Starburst

Star formation is represented using the Siebenmorgen & Krügel (2007) models which assume spherical symmetry and an ISM with dust properties characteristic of the Milky Way. Emission is broken down into two components: an old stellar population uniformly distributed through the volume and hot, luminous O and B stars embedded in dusty hot spots. The density of OB stars is centrally peaked although the model output is the emission integrated over the entire starburst.

The free parameters for the model are: starburst radius r , total luminosity L_{SB} , ratio of luminosity in O and B stars to total luminosity f_{OB} , visual extinction from the center to the edge of the nucleus A_V , and dust density in hot spots around O and B stars n .

Owing to the spherical symmetry of this starburst model and the known ring morphology of the star formation in Cygnus A, parameters such as A_V and the radius r do not have straightforward interpretations here. The assumption of optically thick star formation does not have strong evidence (either for or against), given the absence of high-resolution far-infrared observations.

3.3. Synchrotron

The strong radio emission in Cygnus A justifies the final model component. VLA core fluxes are consistent with power-law emission to the submillimeter where thermal emission begins to dominate. Extrapolating the power law to higher frequencies suggests that the synchrotron spectrum must either be modified or subjected to attenuation. The AGN is known to sit behind a dust lane with significant extinction ($A_V = 50 \pm 30$; Djorgovski et al. 1991). Though some attenuation of the

synchrotron flux is expected at higher frequencies, this alone is not sufficient to explain why the synchrotron power law does not continue through the mid-infrared. With extinction alone, the flux densities at 30 and 10 μm would be similar for the observed spectral index of $\alpha = 0.18$. For $A_V = 50$, the expected flux from the extrapolated synchrotron emission would by itself exceed the measured infrared flux at 10 μm . We conclude that there must be a break in the population of emitting electrons.

One possible explanation is a simple cutoff in the population of relativistic electrons at some energy (Case I). This would manifest itself as a cutoff in the synchrotron spectrum:

$$F_\nu \propto \nu^{-\alpha_1} e^{-\frac{\nu}{\nu_c}} e^{-\tau_r(\nu)}, \quad (1)$$

where ν_c is the frequency corresponding to the cutoff in the energy distribution of the particles, α_1 is the spectral index in the optically thin regime, and $\tau_r(\nu)$ is the dust screen between the synchrotron emitting region and the observer. $\tau_r(\nu)$ follows Draine & Lee (1984).

In Cygnus A the spectral index α_1 is measured from VLA core radio fluxes, and the amplitude of the power law fixed from the same observations. The only free parameters are the cutoff frequency ν_c and the optical depth τ_r . These parameters are partly degenerate in that in trial runs we experienced a situation where τ_r and ν_c would both increase, effectively offsetting each other. To combat this we limited τ_r such that A_V did not exceed 500 toward the radio source.

An alternative model for the synchrotron emission at higher frequencies is a broken power law behind a dust screen (Case II). Aging of the population of relativistic electrons results in a broken power law whose spectral index increases for frequencies higher than a break frequency (Kardashev 1962). The functional form adopted is

$$F_\nu \propto e^{-\tau_r(\nu)} \times \begin{cases} \nu^{-\alpha_1} & \text{if } \nu < \nu_{\text{break}} \\ \nu^{-\alpha_2} & \text{if } \nu \geq \nu_{\text{break}} \end{cases} \quad (2)$$

For a simple aging of the electron population, the post-break spectral index in Cygnus A would be $\alpha_2 = 1.24$. As with Case I, we fit a dust screen in front of the synchrotron emission (τ_r).

3.4. Stellar Contribution to the Mid-infrared

Starlight can potentially contribute to the mid-infrared flux, especially in a large aperture. Using the flux of the stellar component from Jackson et al. (1998) and their brightness profile, we determined the contribution due to starlight in a 20'' aperture. The relative fluxes were consistent with an elliptical galaxy template from Silva et al. (1998). Expected flux densities at 2.2, 5, and 10 μm were computed using the same template spectrum. After subtracting a nuclear point source, the K -band flux density in starlight is consistent with the flux in a 20'' aperture as measured using Two Micron All Sky Survey (2MASS) data products (Skrutskie et al. 2006). At 5 and 10 μm , the starlight contributes roughly 14% and 2%, respectively, of the flux measured by IRS. Therefore, we do not include any contribution to the mid-infrared flux from an old stellar population in our modeling.

3.5. Dust in the NLR

Warm dust in the NLR directly illuminated by the AGN can also contribute to the mid- and far-infrared emission (e.g., Groves et al. 2006). Ramos Almeida et al. (2009) find that a significant contribution to the mid-infrared flux can come from

sources other than the AGN torus, particularly additional hot dust. In Cygnus A Radomski et al. (2002) find an extended component to the mid-infrared emission which is consistent with $T \sim 150$ K dust. Some of this extended emission is cospatial with sites of possible star formation. Dust of this temperature is well reproduced with our choice of starburst models (with the implicit assumption that this dust is heated by star formation).

Higher temperature dust is also present in the inner regions of Cygnus A. However the resolved mid-infrared images of the nuclear regions by Radomski et al. (2002) are unable to distinguish between dust in the ‘torus’ and dust in the NLR heated by the AGN. Without strong observational motivation for an additional dust component we opt not to include one. As will be shown in the next section we are able to reproduce the observed emission without this additional component.

3.6. Prior Constraints on Model Parameters

Prior to fitting the models, the available parameter space was constrained using results from previous studies of Cygnus A. The opening angle and covering fraction of the CLUMPY torus are determined from the σ and N parameters (see Equations (3) and (4) in Mor et al. 2009). We define the half opening angle of the torus as the angle at which the escape probability of a photon drops below e^{-1} :

$$\theta_{\text{half}} = 90 - \sigma \sqrt{\ln N_0}. \quad (3)$$

Tadhunter et al. (1999) measured the opening angle of the ionization cone in Cygnus A using near-infrared *HST* data, finding $\theta_{\text{half}} = (58 \pm 4)^\circ$. We use this value as the torus opening angle and limit the parameter range of σ and N according to Equation (3).

The inclination range of the torus i was limited by VLBI observations and modeling of the inner pc-scale jet ($50^\circ \leq i \leq 85^\circ$; Sorathia et al. 1996). Additionally, the available VLA core radio fluxes were used to fix the synchrotron spectrum at radio frequencies (α_1 and the amplitude).

4. RESULTS

To fit the observed SED, the flux from each model was summed in each wavelength bin and compared to the observed flux within a 20'' aperture. Optimization was performed using Levenberg–Marquardt least-squares minimization.

Some degeneracy between parameters was seen in the model results. Comparison model runs with fewer prior constraints on the parameters resulted in similar χ^2 values to those quoted below for regions of parameter space which are unlikely to be physically reasonable matches to Cygnus A (e.g., CLUMPY torus fits with covering fractions of unity and $\theta_{\text{half}} = 0^\circ$). Prior constraints from other observations are thus important in eliminating regions of parameter space which may be statistically reasonable but physically unrealistic.

After the models were fit, confidence intervals were determined separately for Case I and Case II through bootstrapping with replacement (Efron 1981). These were used to determine the ‘acceptable’ range of parameter values.

The results of the fits are shown in Figure 4. The range of parameter values for fits within at 95.4% confidence interval are shown in Figures 5–9. Section 4.1 discusses the Case I fits (exponential cutoff in the relativistic electron population) and Section 4.2 discusses the Case II fits (broken power law for the synchrotron emission).

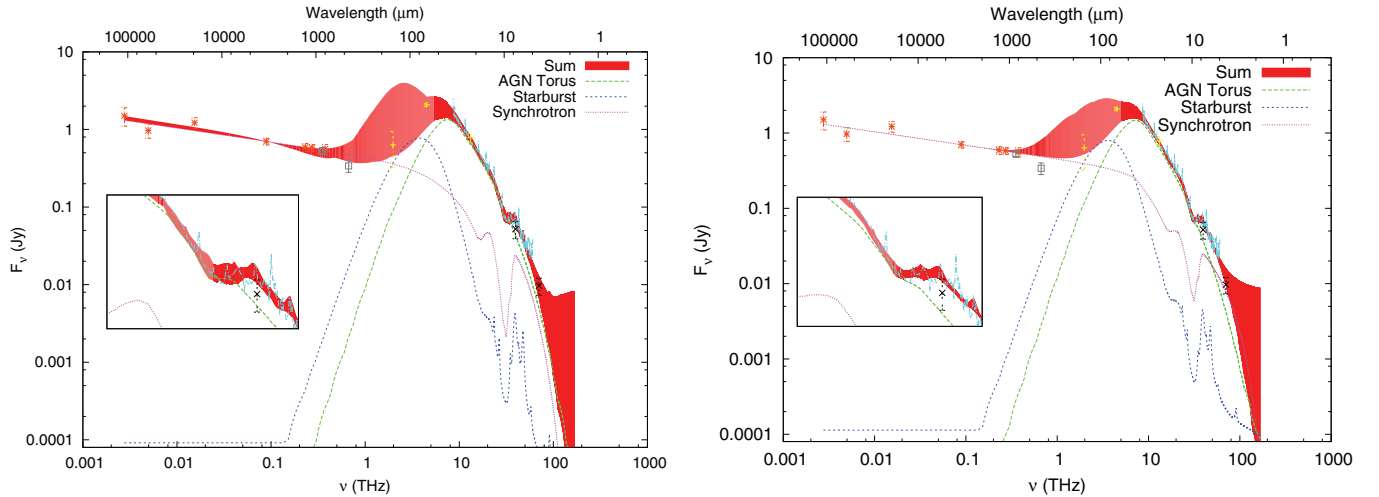


Figure 4. Left: Case I fit to the Cygnus A SED. The shaded region shows all acceptable model fits within the 95.4% confidence interval. The lines are the components for the best-fit (lowest χ^2). Right: same as left, but for Case II fit to the Cygnus A SED. Insets are a zoom of the region around the $9.7 \mu\text{m}$ SiII absorption.

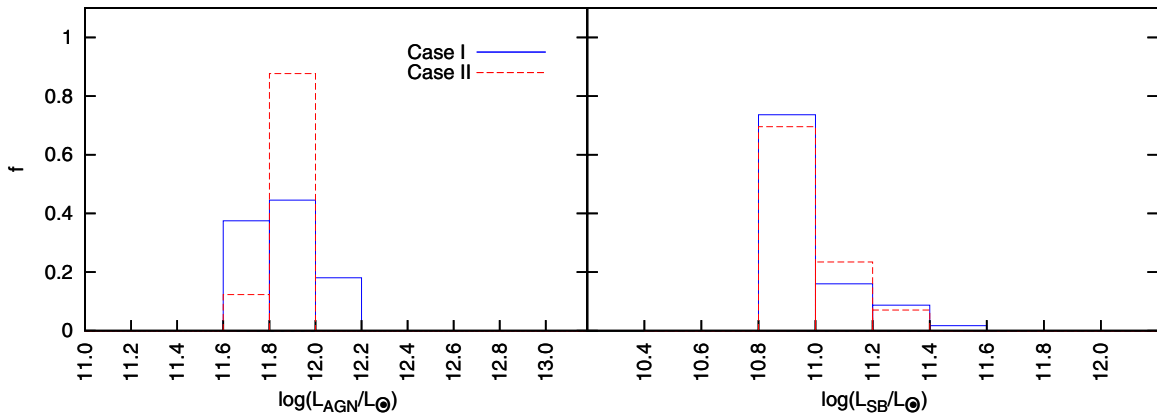


Figure 5. Histogram of bolometric AGN luminosity from best-fit torus parameters (left) and bolometric starburst luminosity (right) for Cases I (blue, solid line) and II (red, dashed line).

4.1. Case I: Synchrotron Exponential Cutoff

The best-fit combination of parameters for Case I has $\chi^2/\text{DOF} = 1.35$. The range of acceptable matches for the 1641 model combinations is shown in Figure 4 (left). Histograms of the best-fit parameters are given in Figures 5 (left) and 6, where f is the fraction of models within the 95.4% confidence interval in a given bin. Parameters for the best-fit model are marked with a small blue horizontal bar.

4.1.1. AGN/Torus Properties and Contribution

The fits favor a bolometric accretion disk luminosity of $\log(L_{\text{AGN}}/L_{\odot}) \sim 11.8\text{--}12.0$ (median of 11.82 with an interquartile spread of 0.09). The fits clearly favor an extended torus ($Y = 200$) with low opacity clouds ($A_V \sim 10\text{--}30$). For the median bolometric luminosity, the inner radius of the torus is $R_d = 0.6$ pc, giving a corresponding outer torus edge of ~ 125 pc (for $Y = 200$).

The range of values for i was limited based on previous work in the radio regime. The fits prefer an inclination for the torus on the high end of the range, $i \approx 80^\circ$. N has a bimodal distribution, with 25% of fits having $N \leq 6$ clouds along an equatorial line of sight and $\sim 60\%$ having $N \geq 20$ clouds. Most of the fits within

the 95.4% confidence interval have a flat radial distribution of clouds ($q = 0$).

The opacity of individual clouds is anti-correlated with the radial extent of the torus; small torus sizes favor high A_V values. High values of N (i.e., many clouds along a given line of sight) are favored for small torus sizes; however, for large torus sizes, both small and large N values are acceptable.

We calculate the total extinction through the torus along both an equatorial line of sight (Figure 6 middle, bottom row), and along our line of sight to the torus (Figure 6 (right), bottom row). The integrated equatorial A_V spans a range between 100 and 250. When viewing angle is taken into consideration, the range narrows, with most fits showing the line of sight A_V between 80 and 120. However the A_V is poorly constrained beyond having a well-defined lower bound.

The torus covering fraction was computed using the method of Mor et al. (2009, their Equations (3) and (4)). By design, the derived covering fractions for the torus are between 50% and 70%.

4.1.2. Starburst

The median value of the starburst luminosity for Case I is $\log(L_{\text{SB}}/L_{\odot}) \sim 10.8$, with a tail up to ~ 11.6 containing approximately 40% of the fits (Figure 5). This covers a range

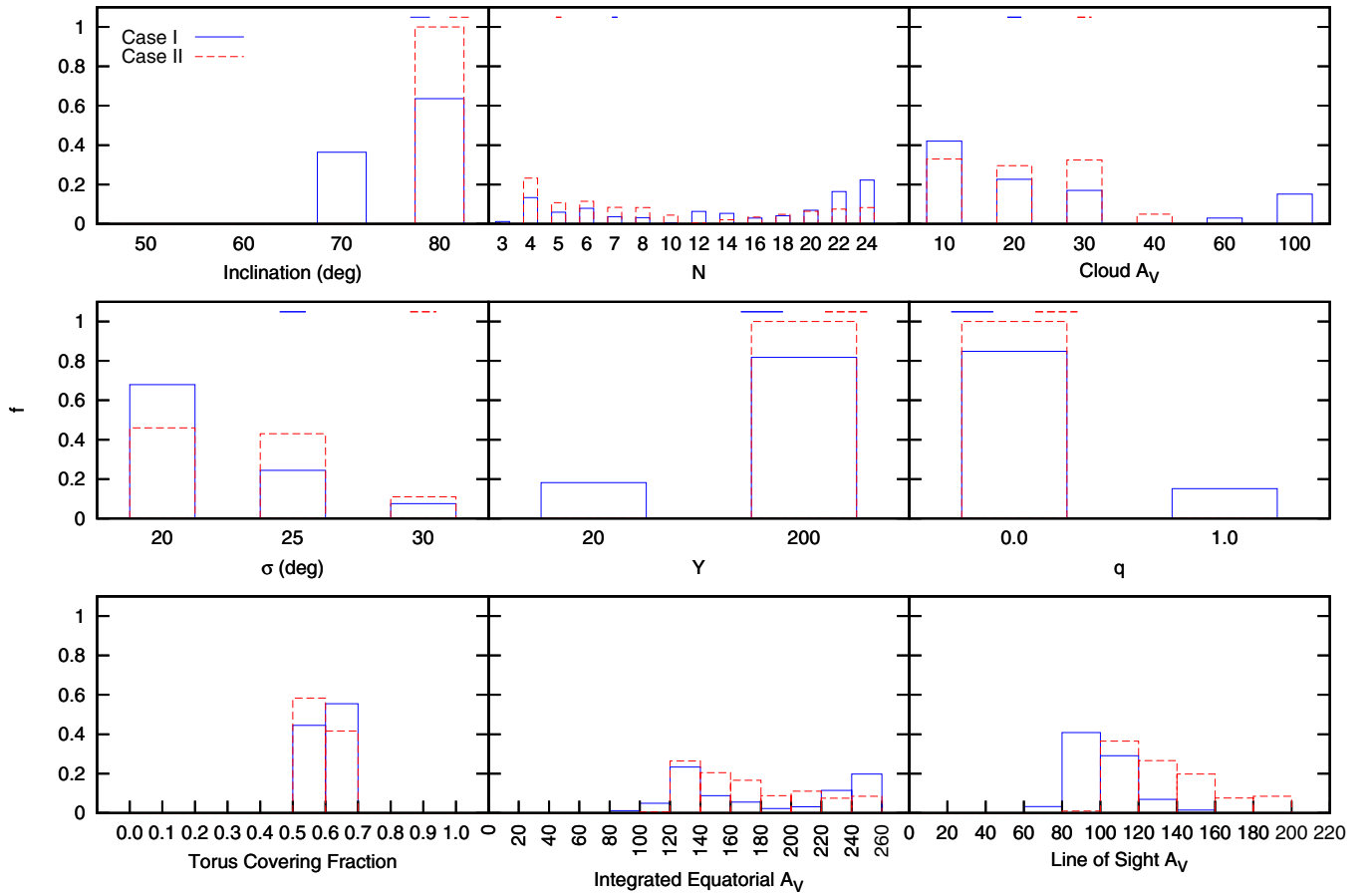


Figure 6. Histograms of torus parameters for acceptable fits within the 95.4% confidence interval for Cases I (blue, solid line) and II (red, dashed line). The small horizontal bars denote the parameter value for the best-fit model. From left to right, top to bottom: viewing angle of the torus (90 meaning the torus is viewed edge-on), N : average number of clouds along an equatorial line of sight, A_V : for an individual cloud, σ : angular width of the distribution of torus clouds, Y : ratio of inner to outer radii, q : power-law index of the radial distribution of torus clouds, covering fraction of the torus (equivalent to the escape probability of optical/UV photons, assuming optically thick clouds), integrated A_V along an equatorial line of sight, integrated A_V along our line of sight (calculated using N , σ , the cloud A_V , and the inclination). Parameters in the first two rows are discrete and the histograms represent the relative number of models with those specific values. Parameters in the bottom row are derived from the parameters in the top two rows as well as the luminosity of the AGN component (shown in Figure 5), and are continuous.

(A color version of this figure is available in the online journal.)

of star formation rates from 10 to 70 $M_\odot \text{ yr}^{-1}$, as determined by the L_{IR} calibration from Kennicutt (1998b) (for a review of SFR estimates, see Kennicutt 1998a):

$$\frac{\text{SFR}}{(M_\odot \text{ yr}^{-1})} = 4.5 \times 10^{-44} \frac{L_{\text{FIR}}}{\text{erg s}^{-1}} = 1.72 \times 10^{-10} \frac{L_{\text{FIR}}}{L_\odot}. \quad (4)$$

Roughly half the fits within this confidence interval have 40% of the starburst luminosity in the form of OB stars. The distribution of size parameters is relatively flat, although even the largest sizes would be unresolved by our *Spitzer* observations. The distribution of dust density n peaks around 10^3 cm^{-3} . The extinction through this starburst is relatively unconstrained.

The strengths of dust features in starburst models falling within the 95.4% confidence interval were also measured using PAHFIT to compare with the intensities of observed dust feature. Figure 8 shows histograms of the predicted intensities from our models divided by the measured intensity from the IRS spectrum. Only dust features which are detected in the IRS spectrum have been plotted. In general the agreement is good with most models matching the measured line intensities. However for some dust features a significant number of the models within the confidence interval predict emission in excess

of what is observed. While a detailed comparison of the relative strengths of dust features is beyond the scope of the paper, we suggest the models are able to generally reproduce the dust features seen in the spectrum. The dust features which show the greatest discrepancy between observed fluxes and those predicted from the modeling (6.2, 7.8, and 11.3 μm) comprise three of the five lowest signal-to-noise dust feature fits in the spectrum. The discrepancy may then be due to the difficulty of fitting low equivalent width dust features.

PAHFIT also attempted to fit eight other dust features in the IRS spectrum. The upper limits for three (5.7, 14.0, and 15.9 μm) are consistent with expectations from the starburst models. One dust feature (16.4 μm) has a measured limit below that expected from the best-fitting starburst model, and so is discrepant. The other four features (7.6, 10.7, 11.2, and 12.7 μm) also have upper limits from PAHFIT which are lower than the expected value from the starburst models. However, these are coincident with or near other spectral features (e.g., narrow emission lines or the SiI absorption). Thus an accurate measurement of these dust features in the IRS spectrum would rely more heavily on fitting the wings, which could prove difficult in a continuum dominated source such as Cygnus A.

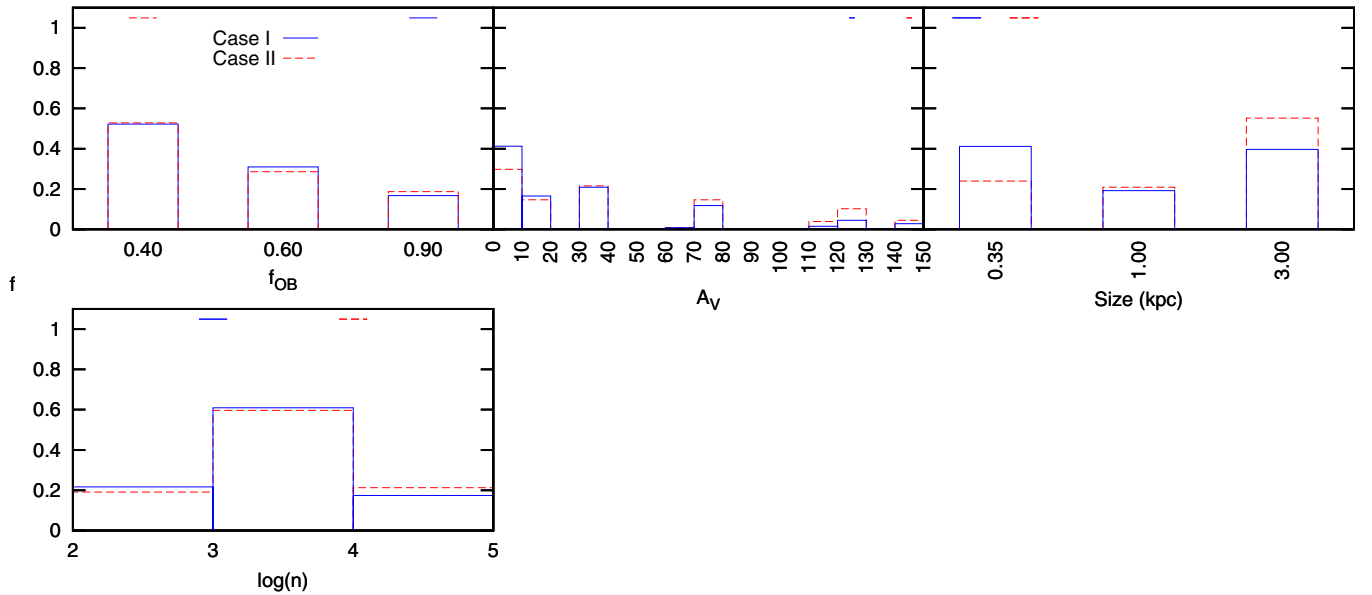


Figure 7. Histograms of starburst parameters for acceptable fits within the 95.4% confidence interval for Cases I (blue, solid line) and II (red, dashed line). The small horizontal bars denote the parameter value for the best-fit model. From left to right, top to bottom: the fraction of luminosity in O and B stars, integrated extinction through the starburst, size of the starburst, dust density in hot spots around O and B stars. f_{OB} and the size are discrete and the histograms represent the relative frequency of models with those specific values. The A_V and n parameters are more finely sampled than the histogram represents; the plots show the relative fraction of models with parameters in the noted ranges.

(A color version of this figure is available in the online journal.)

4.1.3. Synchrotron Radiation

The synchrotron emission amplitude was fixed using the non-thermal emission from the radio core, assuming it to be a point source at all frequencies observed. The synchrotron model, therefore, has only two parameters: cutoff frequency (ν_c) and extinction due to dust (τ_r).

The distribution of cutoff frequencies is somewhat broad, covering the range of $\nu \approx 10\text{--}60$ THz ($5\text{--}30\ \mu\text{m}$) (Figure 9). The fits show a range of acceptable extinction for the dust screen, peaking in the range of $A_V \approx 60\text{--}80$, slightly lower than the predicted line of sight A_V from the torus models. The integrated luminosity of the “core” synchrotron flux is $\log(L_{\text{sync}}/L_\odot) \sim 11.2$.

This break in the synchrotron spectrum is consistent with other powerful FR II radio sources, where an extrapolation of the radio synchrotron emission significantly exceeds the observed optical flux (e.g., Schwartz et al. 2000; Sambruna et al. 2004; Mehta et al. 2009). In these cases, the similarity of the radio and X-ray spectral indices suggests that the X-rays may be produced by inverse Compton (IC) scattering.

With the modeled synchrotron spectrum, a magnetic field strength, and the assumption that each electron emits at a single frequency, the energy distribution of relativistic electrons can be determined:

$$\gamma(\nu) = \sqrt{\frac{4\pi m_e c \nu}{3eB}}, \quad (5)$$

where all constants and values are in the cgs system of units.

The magnetic field strength has been estimated from VLBI observations to be roughly 17 mG (100 mG) in the jet (core) (Roland et al. 1988; Kellermann et al. 1981). Using $\nu \approx 30$ THz and Equation (5) we find $\gamma \sim 2 \times 10^4$ (8×10^3) for the jet (core).

In addition to radiating energy via synchrotron emission, relativistic electrons can also lose energy via IC scattering. The ratio of the energy lost through synchrotron to the energy lost to

IC is simply given by the ratio of the magnetic and radiation energy densities, with the cosmic microwave background (CMB) setting a minimum value for the radiation energy density. Using the magnetic field strength assumed above (17 mG), and $T_{\text{CMB}} = 2.73$ K, synchrotron losses are dominant by a factor $> 10^8$. A larger magnetic field would further enhance this ratio, so IC losses from scattering of CMB photons have a negligible effect on the population of relativistic particles.

4.2. Case II: Synchrotron Broken Power Law

The range of best-fit models for Case II is shown in Figure 4 (right). The best-fit model had $\chi^2/\text{DOF} = 1.03$. Figure 4 (right) shows the range of the 1580 model fits within the 95.4% confidence interval. Histograms of the best-fit parameters are given in Figures 5 (left) and 6. A red horizontal bar marks the parameter for the parameters with the lowest χ^2 value.

4.2.1. AGN/Torus

In Case II fits the bolometric accretion disk luminosity is roughly $\log(L_{\text{AGN}}/L_\odot) \sim 11.8$ (median 11.88 with an interquartile range of 0.08). The distribution of luminosities is more strongly peaked than for Case I, with over 80% of fits occupying the $\log(L/L_\odot) = 11.8\text{--}12.0$ bin. A large torus ($Y = 200$) is exclusively preferred, corresponding to $R_{\text{out}} \approx 135$ pc (using the median value for the luminosity). Again, by design, the torus covering fraction is between 0.5 and 0.7. An inclination of $i = 80^\circ$ is exclusively preferred.

The distribution of the average number of clumps along a line of sight N is again bimodal with either large (<14) or small (<8) numbers of clouds preferred. The number of clouds is anticorrelated with the extinction through individual clouds. The spread in equatorial A_V is somewhat large (120–260), but the line of sight A_V is more constrained (100–160 for 80% of fits).

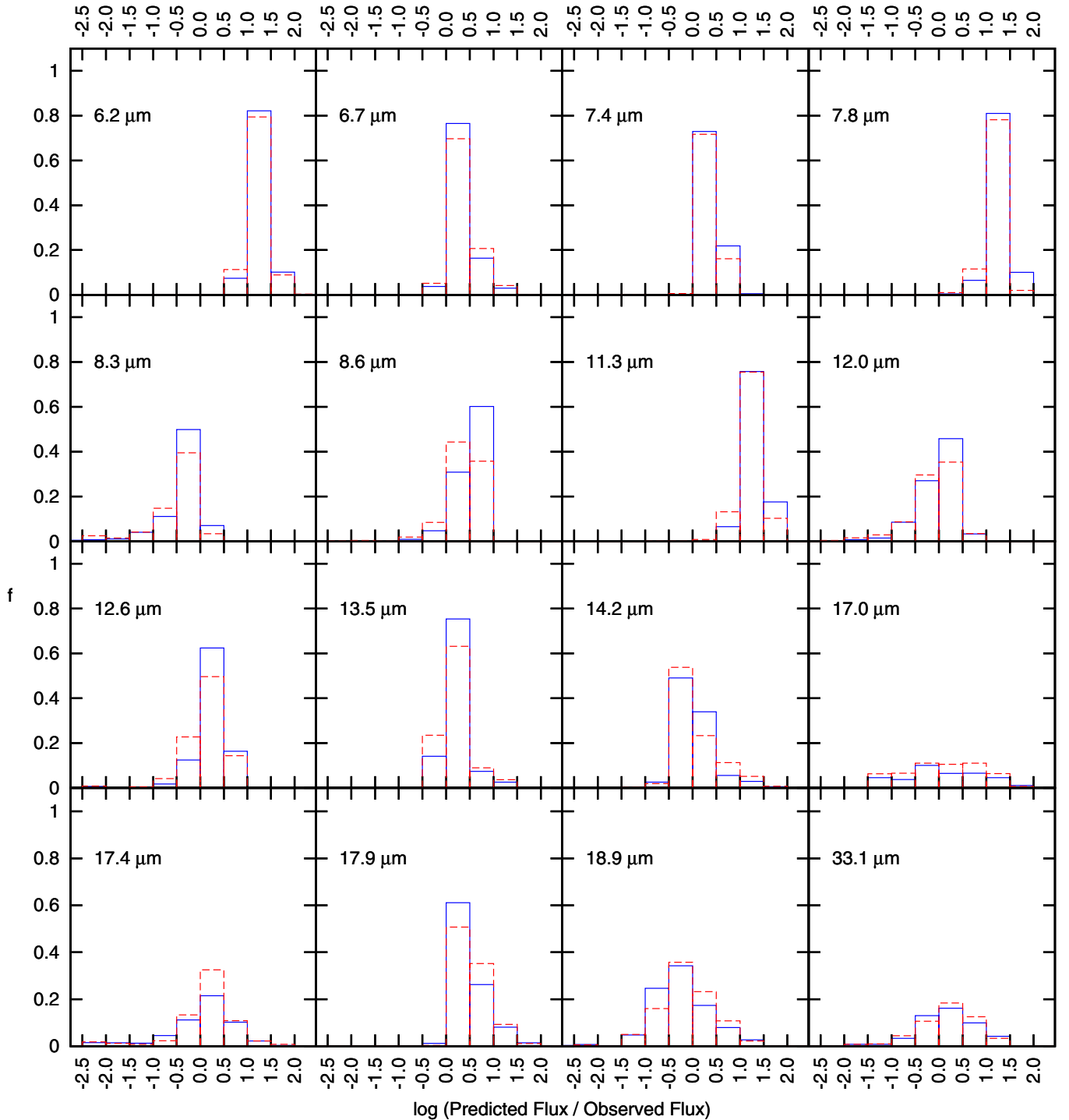


Figure 8. Histograms of $\log(\text{predicted flux}/\text{measured flux})$ for dust features detected in the IRS spectrum of Cygnus A. The predicted values were measured from the Siebenmorgen & Krügel (2007) models using PAHFIT. The solid blue line denotes Case I while the red dashed line denotes Case II (see Table 2 for the measured flux values).

(A color version of this figure is available in the online journal.)

4.2.2. Starburst

The starburst component in Case II fits show qualitatively similar behavior to Case I. The typical luminosity is comparable, but shows a smaller tail up to $\log(L_{\text{SB}}/L_{\odot}) = 11.6$ ($\sim 30\%$ of fits). Other parameters are similar to those in Case I fits, with the exception of the size, where there appears to be a very slight preference for a larger size, suggesting an overall cooler dust temperature for a given luminosity. This may indicate

a degeneracy between the starburst and synchrotron models. The critical frequency for the synchrotron spectrum (see the next section) is at lower frequency for Case II, resulting in a smaller contribution to the far-infrared flux. The starburst model compensates with a larger overall size to provide additional cool dust emission (for a given starburst luminosity). The dust feature strengths expressed in the models generally compare favorably with Case I.

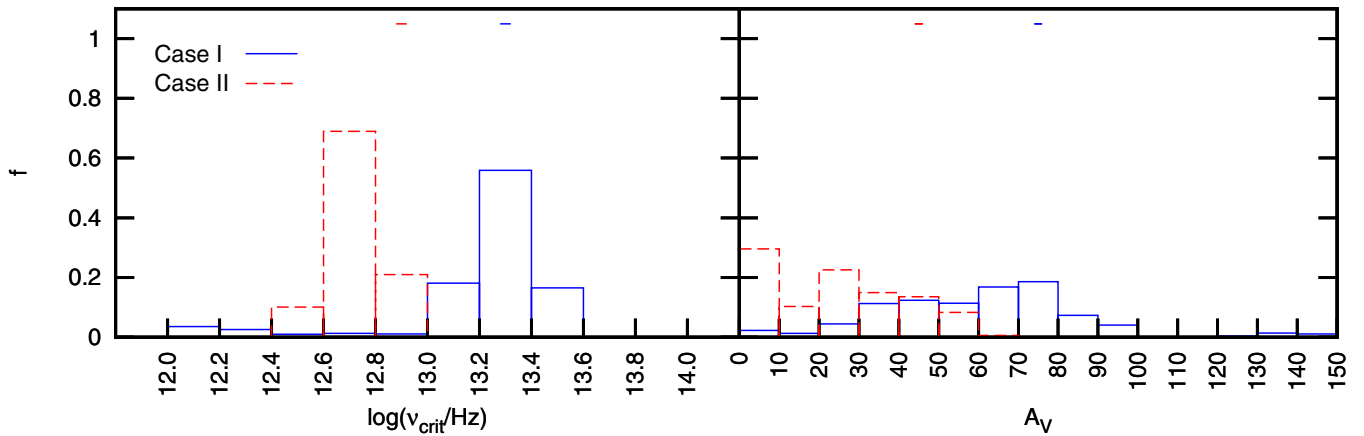


Figure 9. Histograms of synchrotron parameters for acceptable fits within the 95.4% confidence interval for Cases I and II. For Case I $\nu_{\text{crit}} \equiv \nu_{\text{cutoff}}$ and for Case II $\nu_{\text{crit}} \equiv \nu_{\text{break}}$. $A_V = 1.08\tau_r$ ($\nu = V$ -Band).

(A color version of this figure is available in the online journal.)

4.2.3. Synchrotron Properties and Contribution

An alternative mechanism for limiting the influence of synchrotron emission at shorter wavelengths is for the spectrum to break at some frequency (see Section 3.3). Case II fits used a fixed pre-break spectral index of $\alpha_1 = 0.18$ and a post-break spectral index of $\alpha_2 = 1.24$, consistent with aging of the relativistic population (without injection of additional particles; Kardashev 1962). As the flux density decreases in a slower fashion when compared to an exponential cutoff, the power law must break at lower frequencies to ensure that the observed mid-infrared flux is not exceeded.

The break frequency is roughly 5 THz (60 μm). The unobscured synchrotron luminosity is found to be $\log(L_{\text{sync}}/L_{\odot}) \sim 11.1$. Following the same arguments as Case I, the electrons emitting at the break frequency have $\gamma \sim 7 \times 10^3$ (3×10^3) for jet (core) magnetic field values. Again, this is consistent with the observed properties of jets in other FR II radio sources.

The extinction of the dust screen in front of the radio source is lower than Case I fits, with $A_V < 60$, still within range of the extinction to the central source estimated by Djorgovski et al. (1991), but lower than the computed line of sight A_V from the torus. The larger discrepancy compared with Case I is due to the enhanced short wavelength emission of the broken power law at shorter wavelengths (when compared to the Case I exponential cutoff). The equatorial torus A_V increases to provide an overall cooler temperature, and the modeled dust screen in front of the synchrotron component remains small (in order to contribute sufficient flux at shorter wavelengths).

5. DISCUSSION OF GENERAL SED RESULTS

5.1. Luminosity and Kinetic Power in Cygnus A

Our best estimate of the bolometric AGN luminosity in Cygnus A ($\log(L/L_{\odot}) \sim 12$, including the synchrotron component and X-ray emission) is above the Whysong & Antonucci (2004) estimate using Keck mid-infrared observations and a PG-quasar spectrum ($3.9 \times 10^{11} L_{\odot}$), but is somewhat below the estimates of Tadhunter et al. (2003) who find $L_{\text{bol}} = (12\text{--}55) \times 10^{11} L_{\odot}$ (although consistent with the low end of their range). The bolometric AGN luminosity is the best constrained parameter and is insensitive to the synchrotron model adopted.

The kinetic power in the expansion of the radio lobes in Cygnus A can be inferred from X-ray observations of the cluster environment (Wilson et al. 2006). From this analysis of *Chandra* observations Cygnus A has a kinetic power of $L_{\text{kin}} = 1.2 \times 10^{46} \text{ erg s}^{-1}$ ($\log(L_{\text{kin}}/L_{\odot}) = 12.5$), a factor of three larger than the bolometric AGN luminosity inferred from the modeling. Given the uncertainties in the determinations of both the kinetic and bolometric power, it is unclear if the kinetic power dominates significantly over the power emitted as radiation.

5.2. Implications of the Modeling

The probable torus sizes from the SED modeling give outer radii of $R_o = 130 \text{ pc}$ (≈ 0.2 at the distance of Cygnus A). Generally, the torus parameters provide reasonable estimates of the properties of the obscuring structure. This predicted outer radius is significantly larger than what is generally assumed to be reasonable for the obscuring torus. A torus with an angular diameter of ≈ 0.2 is well within the range of multiple mid- and near-infrared instruments. At large radii, the distinction between torus clouds and NLR clouds may be somewhat arbitrary. The observations by Canalizo et al. (2003) show extended emission on scales including and larger than our fit torus sizes. The degree of contamination by emission lines is unclear, and it would be interesting to attempt to obtain a line-free continuum image to ascertain the true size and shape of the continuum emitting region.

The torus size parameters are consistent with those found from modeling of the $9.7 \mu\text{m}$ Sil feature by Imanishi & Ueno (2000). Their radiative transfer modeling suggested some torus properties similar to those presented here, namely, a small inner radius ($< 10 \text{ pc}$) and an inner-to-outer radius of 80–500. In contrast they find a steeper radial dependence for the dust distribution $q \sim 2\text{--}2.5$, possibly due to the use of a smooth dust distribution.

The preferred parameters for the obscuring torus; low q , high A_V , and large Y suggest the model is being driven toward a cooler spectrum. This could influence the starburst component, leading to a lower inferred star formation rate. In order to combat this it would be beneficial to place tighter constraints on the far-infrared SED to enable better modeling of the star formation in Cygnus A.

A comparison of our IRS data with mid-infrared observations by Radomski et al. (2002) is consistent with the suggestion of our model that at 10 and 18 μm , the emission is dominated by the torus + synchrotron component. Tadhunter et al. (1999) measure a K -band nuclear point source with $F_\nu = (49 \pm 10) \mu\text{Jy}$. This flux limit is broadly consistent with some model fits for the torus + synchrotron component.

For the estimated bolometric luminosity $10^{12} L_\odot$ and a black hole mass of $2.5 \times 10^9 M_\odot$, the Eddington ratio (L/L_{edd}) is $\sim 1.3 \times 10^{-2}$. This is similar to, but slightly lower than, previous estimates from Tadhunter et al. (2003), likely due to the fact that our model attributes some of the IR luminosity to the heating of dust from star formation. This starburst luminosity is relatively well constrained, with consistent values across the bulk of fits. The AGN (torus plus jet) contributes $\sim 90\%$ of the infrared luminosity and star formation produces the remaining 10%.

5.3. Future Work and Observations

The results of our modeling can be tested and improved with the help of future observations in various wavelength regimes.

Around 1 THz ($\sim 300 \mu\text{m}$) the synchrotron and starburst model components are of similar flux density, with the synchrotron contribution decreasing and the thermal contribution from star formation beginning to dominate. Unfortunately, data at this location are limited in resolution, sensitivity, and wavelength coverage. Disentangling the emission from the starburst and emission from the AGN at this frequency will be possible with the Atacama Large Millimeter Array (ALMA; e.g., Carilli 2005), particularly with the availability of Band 10 and full science capabilities.

The far-infrared suffers from poor sampling of the SED. In this regime the *Herschel Space Observatory* (Pilbratt et al. 2010) provides an opportunity to improve on the understanding of the continuum emission. SPIRE and PACS cover wavelength ranges of interest: the region of the possible jet-break and the long wavelength side of the infrared bump. Observations here can provide improved constraints for the synchrotron and starburst models. González-Alfonso et al. (2010) demonstrate decomposition of far-infrared emission for Mrk 231, with the *Herschel* SPIRE observations providing important constraints. Similar observations of Cygnus A will provide important data on this sparsely observed portion of the SED. PACS observations would contribute measurements of far-infrared fine structure lines which could be used in concert with the mid-infrared lines to provide an alternative method of determining the relative contribution of star formation and AGN activity to the infrared luminosity (e.g., Fischer et al. 2010, for Mrk 231).

Although the synchrotron spectrum breaks in our fits, it may still dominate the flux between 5 and 10 μm . This result can be tested by future mid-infrared polarimetry or variability studies. Synchrotron emission from a compact source such as the radio core or jet knots should be subject to flaring. If the emission in this wavelength range is instead of a thermal origin (e.g., hot dust in the host galaxy), the 10 μm flux will remain relatively stable.

6. SUMMARY

Using a combination of a new mid-infrared spectrum from the *Spitzer Space Telescope* and radio data from the literature, the nuclear emission in Cygnus A has been modeled as a combination of power-law emission from a synchrotron jet,

reprocessed AGN emission from a dusty torus, and emission from a dusty circumnuclear starburst.

The data are well fit by a combination of these three models, and all three are necessary to reproduce the observed emission. Statistically acceptable fits were found for both an exponential cutoff in the population of relativistic electrons (Case I) as well as emission from an aging electron population (Case II); however we are unable to distinguish between the two cases. For Case I we find the cutoff frequency to be between 10 and 50 THz (5–30 μm) while for Case II fits the predicted break frequency is 5 THz (60 μm). Degeneracy between the starburst and synchrotron components makes a more precise determination of the break or cutoff frequency difficult. Better observations on the long-wavelength side of the thermal bump will provide tighter constraints on the models.

From this modeling, we find the following.

1. The bolometric luminosity of the AGN in Cygnus A is $\sim 10^{12} L_\odot$.
2. The mid-infrared emission is consistent with emission from a clumpy obscuring torus with an outer size of ~ 130 pc.
3. The far-infrared emission is consistent with being dominated by star formation which is occurring at a rate between 10 and 70 $M_\odot \text{ yr}^{-1}$.

In Cygnus A, the infrared emission is a combination of AGN and starburst heated dust, with the AGN contributing $\sim 90\%$ of the luminosity.

The authors thank the anonymous referee whose comments have improved the quality of this paper.

This work is based on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA. Support for this work was provided by NASA through an award issued by JPL/Caltech. Computing resources were provided in part by the Research Computing group of the Rochester Institute of Technology. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. This research has made use of NASA's Astrophysics Data System. This work is supported in part by the Radcliffe Institute for Advanced Study at Harvard University.

G.P. thanks D. Whelan, D. M. Whittle, and A. Evans for helpful discussions.

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